



## *Planet Detection Metrics:*

Window and One-Sigma Depth Functions  
for Data Release 25

KSCI-19101-001

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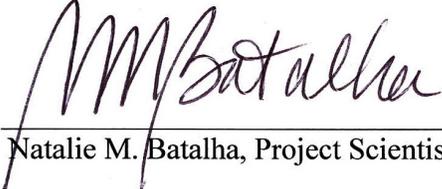
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**Table of Contents**

1. Introduction.....	6
2. Window Function.....	7
3. One-Sigma Depth Function .....	9
4. Window Function Changes and Verification.....	11
5. References.....	13

## 1. Introduction

This document describes the window and one-sigma depth functions relevant to the Transiting Planet Search (TPS) algorithm in the *Kepler* pipeline (Jenkins 2002; Jenkins et al. 2017). The window function specifies the fraction of unique orbital ephemeris epochs over which three transits are observable as a function of orbital period. In this context, the epoch and orbital period, together, comprise the ephemeris of an orbiting companion, and ephemerides with the same period are considered equivalent if their epochs differ by an integer multiple of the period. The one-sigma depth function specifies the depth of a signal (in ppm) for a given light curve that results in a one-sigma detection of a transit signature as a function of orbital period when averaged over all unique orbital ephemerides.

These planet detection metrics quantify the ability of TPS to detect a transiting planet signature on a star-by-star basis. They are uniquely applicable to a specific *Kepler* data release, since they are dependent on the details of the light curves searched and the functionality of the TPS algorithm used to perform the search. This document describes the window and one-sigma depth functions relevant to *Kepler* Data Release 25 (DR25), where the data were processed (Thompson et al. 2016) and searched (Twicken et al. 2016) with the SOC 9.3 pipeline. In Section 4, we describe significant differences from those reported in *Kepler* Data Release 24 (Burke & Seader 2015) and document our verification method.

These data products are archived at the NExScI Exoplanet Archive<sup>1</sup> as two separate FITS files per target as described below. They are available for all targets searched for planets by the *Kepler* pipeline (Twicken et al. 2016) and can be downloaded as a set, or for individual targets through the interactive interface.

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<sup>1</sup> <http://exoplanetarchive.ipac.caltech.edu/index.html>

## 2. Window Function

The window function for a transit survey specifies the fraction of unique orbital ephemeris epochs that permits  $\geq k$  transits to be observed as a function of period for a fixed-length data set (Gaudi 2000; Burke et al. 2006; Burke & McCullough 2014). In this context, ephemerides with the same period are considered identical if their epochs differ by an integer multiple of the period. The *Kepler* pipeline requires  $k \geq 3$  transits for planet detection at each of the fourteen transit durations searched (*i.e.*, 1.5, 2.0, 2.5, 3.0, 3.5, 4.5, 5.0, 6.0, 7.5, 9.0, 10.5, 12.0, 12.5, and 15.0 hrs), so the reported window functions are based on this three-transit criteria. An additional check is performed by TPS when exactly three transits contribute to a putative detection. In the three-transit case only, each transit is inspected. If less than half of the in-transit cadences have full weight (*e.g.*, full weight,  $w=1.0$ , and full deweighting,  $w=0.0$ ) for any of the three transits that contribute to the detection, then the detection is rejected and does not become a Threshold Crossing Event (TCE). The three-transit weight check was not included in the DR24 calculation of the window function (Burke & Seader 2015). In Section 4, we discuss the impact of including this additional check.

The FITS files are named `kplr<identifier>_<release>_window.fits`, where `<identifier>` corresponds to the target's *Kepler* Input Catalog (KIC) identification number and `<release>` specifies the *Kepler* data release. Each file contains 14 binary table extensions for the 14 pulse durations searched by TPS. The transit duration is specified by the keyword TDUR in the FITS header for each extension. Each binary table extension has two columns corresponding to the orbital period (in days) and the window function (unitless) for the stated transit duration. In order to evaluate the window function at an arbitrary orbital period, it is up to the end user to interpolate between the tabulated orbital periods. However, in practice the resolution in orbital period provided is sufficient for 'nearest neighbor' interpolation in most use cases.

Due to the expected increase in transit duration with orbital period for transiting planet signatures, the entire period range is not searched at all transit durations. The shortest transit durations have an upper limit to the orbital periods searched, and the longest transit durations have a lower limit to the orbital periods searched. Thus, the range of orbital period over which valid window functions are computed by TPS varies with duration. Section 2.2 of Tenenbaum et al. (2014) describes details of the orbital period range searched as a function of transit duration. The orbital period range accommodates the expected transit durations assuming circular orbits with extended coverage to allow for moderately eccentric orbits.

Several example window functions are provided in the left-hand panels of Figures 1 – 3. These examples illustrate targets that have varying amounts of observational data available. The window function ranges from 0 to 1. It has a value of one at short periods, where there is a high probability of having at least three transits. The window function then decays to zero at long orbital periods, since there is a decreasing probability that the flux time series is long enough to allow at least three transits to be observed. For nearly all *Kepler* targets, the window function varies insignificantly with transit duration,

so it is typically valid to adopt the results of any small set of durations that spans the full range of periods. This is well illustrated in the left-hand panels of Figures 1 – 3, where all 14 durations are plotted, but only one continuous curve is apparent.

### 3. One-Sigma Depth Function

The detectability of a transit signal is predominately determined by its signal-to-noise ratio (SNR) (Gaudi 2000). One can estimate the transit SNR as  $\Delta\sqrt{N_{\text{tran}}}/\sigma_{\text{cdpp}}$ , where  $\Delta$  is the expected transit depth,  $N_{\text{tran}}$  is the expected number of transits, and  $\sigma_{\text{cdpp}}$  is an averaged noise estimate from the root-mean-square (rms) CDPP (Christiansen et al. 2012). However, this SNR estimate assumes all observations have the same noise properties and that missing or de-weighted data in the flux time series are evenly distributed in time. For the one-sigma depth function, TPS quantifies the transit signal depth that yields a SNR of unity as a function of the orbital period by taking into account the full details of the time varying properties of the noise and accounting for any missing or de-weighted data. When using the one-sigma depth function,  $\Delta_{\text{one}}$ , the transit signal-to-noise ratio is simply  $\text{SNR} = \Delta/\Delta_{\text{one}}$ .

As for the window function, the one-sigma depth function is archived in FITS format. Each file contains 14 binary table extensions giving results for the 14 transit durations searched by TPS. Each binary table extension has two columns corresponding to the orbital period (in days) and the one-sigma depth (in ppm). The file names have the form `kplr<identifier>_<release>_onesigdepth.fits`, where `<identifier>` and `<release>` have the same definitions as for the window function files. In order to evaluate the one-sigma depth function at an arbitrary orbital period, it is up to the end user to interpolate between the tabulated orbital periods. However, in practice the resolution in orbital period provided is sufficient for ‘nearest neighbor’ interpolation in most use cases.

Several examples of one-sigma depth functions are provided in the right-hand panels of Figures 1 - 3. These examples feature targets with varying amounts of missing data to illustrate the impact of incompleteness. In the large  $N_{\text{tran}}$  (*i.e.*, short-period) regime,  $\Delta_{\text{one}}$

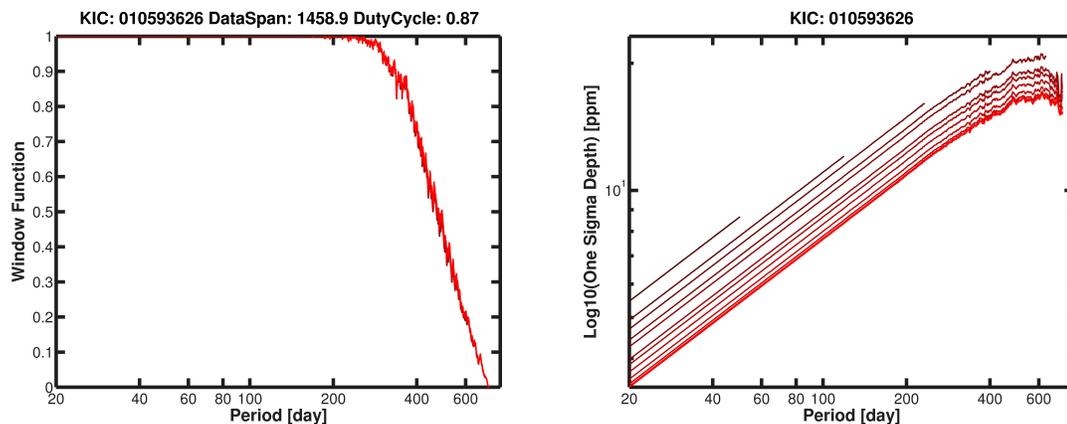


Figure 1. Example planet detection metrics for a target with data available in every quarter (Q1-Q17). This target corresponds to KIC 10593626, the host of *Kepler-22b* (Borucki et al. 2012). **Left:** Window function as a function of orbital period. **Right:** One-sigma depth function as a function of orbital period. The top (bottom) line corresponds to the shortest (longest) transit duration searched for transit signatures.

increases as  $\sqrt{P_{\text{orb}}}$ , where  $P_{\text{orb}}$  is the orbital period. For small values of  $N_{\text{tran}}$  (*i.e.*, long periods),  $\Delta_{\text{one}}$  deviates from this simple dependence on  $P_{\text{orb}}$  because the noise properties of the full flux time series are not well represented by such sparse sampling.

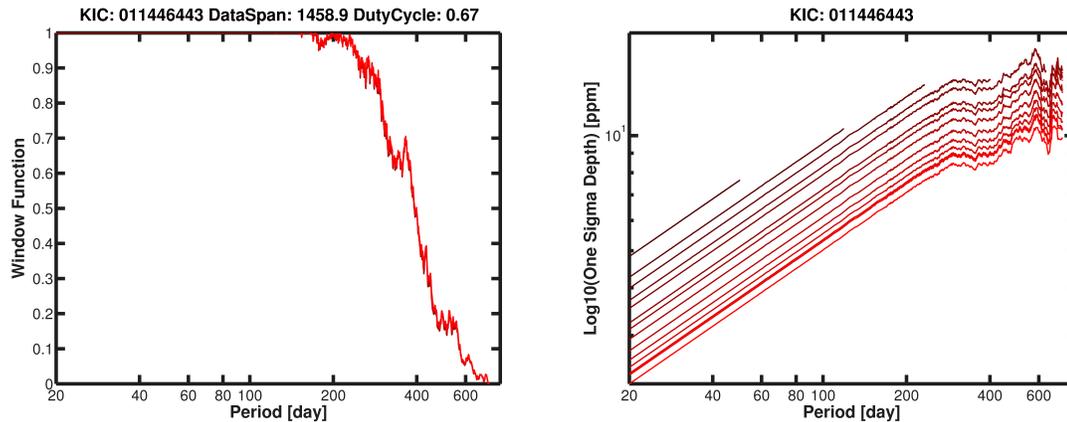


Figure 2. Example planet detection metrics for a target impacted by the loss of Module 3 (*i.e.*, every fourth quarter of data is missing). This target corresponds to KIC 11446443, the host of TrES-2b (O’Donovan et al. 2006). **Left:** Window function as a function of orbital period. **Right:** One-sigma depth function as a function of orbital period. The top (bottom) line corresponds to the (shortest) longest transit duration searched for transit signatures.

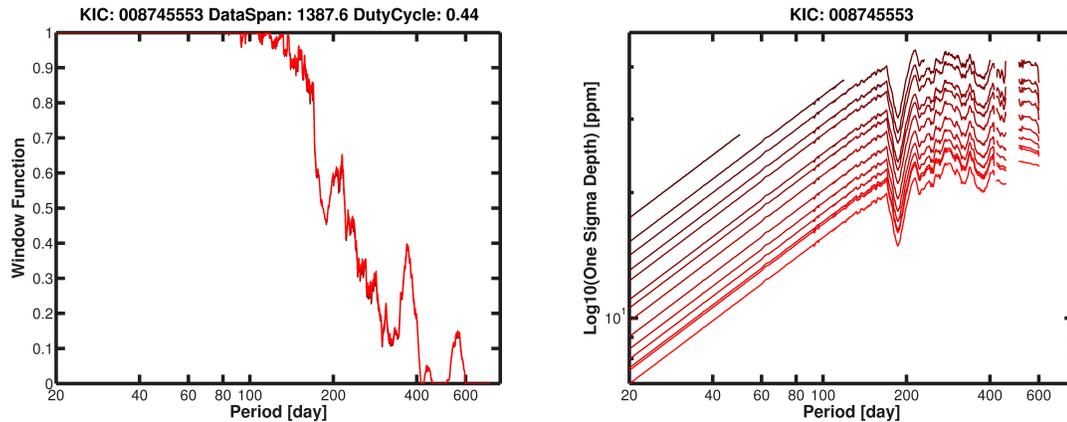


Figure 3. Example planet detection metrics for a target with data missing every other quarter. This target corresponds to KIC 8745553. In general, targets with significant amounts of missing data or patterns in the missing data segments have increased structure in the planet detection metrics as data coherence rapidly changes with orbital period. **Left:** Window function as a function of orbital period. **Right:** One-sigma depth function as a function of orbital period. The top (bottom) line corresponds to the (shortest) longest transit duration searched for transit signatures.

#### 4. Window Function Changes and Verification

In this section, we briefly describe the method used to calculate the window functions and their verification using an alternative Monte-Carlo sampling method. For *Kepler*, the window function requires evaluating, for a given ephemeris (*i.e.*, period and epoch) whether enough data exists to observe three independent transit events and whether the in-transit data passes the three-transit weight check described in Section 2. The results depend on the details of the TPS version employed. Twicken et al. (2016) and Jenkins et al. (2017) describe the SOC 9.3 pipeline employed for DR25.

In principle, the window function could be computed by TPS during the actual transit search within the pipeline. However, in earlier versions, the requisite information was not retained between evaluations at different ephemerides. Therefore, the signal search function in TPS was augmented post facto to accumulate this information across all ephemeris evaluations. With the penalty of a few storage variables, the window function was then calculated outside the formal pipeline with this modified version of TPS.

In parallel with this effort, we developed a Monte-Carlo transit injection and recovery test for TPS known as flux-level transit injection (FLTI). This FLTI uses a highly scaled-down version of TPS that only searches the ephemeris grid points enclosing the injected ephemeris location, rather than the entire period grid of a normal search. In brief, this FLTI code adds realistic limb-darkened transit signals to the flux time series that was previously quarter-stitched and detrended during the full TPS run (see Thompson 2016). The FLTI is based on TPS alone and is completely independent of the other pipeline modules (Jenkins et al. 2017). Hence, this is very different (and many orders of magnitude faster) than the pixel-level transit injection study of Christiansen et al. (2016), which injects transit signals into the calibrated (CAL) pixel-level data and runs the *Kepler* pipeline to perform the photometry (PA), systematic error correction (PDC), full transit search (TPS), and signal evaluation (DV). These shortcuts allow us to test the recovery of hundreds of thousands of realistic transit signatures injected at different ephemerides into the light curve of a single star. In this case, the FLTI code is used to independently calculate an empirical window function that can be used to verify the stand-alone, augmented TPS module's window function calculation. This verification is only practical for a small set of stars because FLTI is orders of magnitude more costly due to the large number of injections required to infer the window function. However, once verified, the augmented TPS code can readily calculate accurate window functions for all *Kepler* targets.

Figure 4 shows the results of the initial verification exercise, comparing the FLTI and augmented TPS window functions for four different stars. The initial verification identified differences as large as 50% between the FLTI window functions (black lines) and the initial augmented TPS window functions (red lines). This discrepancy was traced to the fact that the three-transit weight check was missing from the augmented TPS window function code (*i.e.*, the code only accounted for the three-transit requirement). With the addition of an extra array to retain the results of the weight checks for three-transit events, the final augmented TPS window functions for DR25 (blue lines) are

consistent with the empirical window functions (black lines) computed with FLTI to within the sampling errors of these Monte-Carlo results.

As a cautionary note, the three-transit weight check was also missing from the augmented TPS code used to produce the DR24 window functions (Burke & Seader 2015). At that time, the FLTI code was not available for verification, so this discrepancy was not identified. Consequently, the archival window functions for DR24 are overly optimistic representations of the actual SOC 9.2 pipeline performance. End users of the DR24 window functions are encouraged to consider how this discrepancy might impact their scientific results.

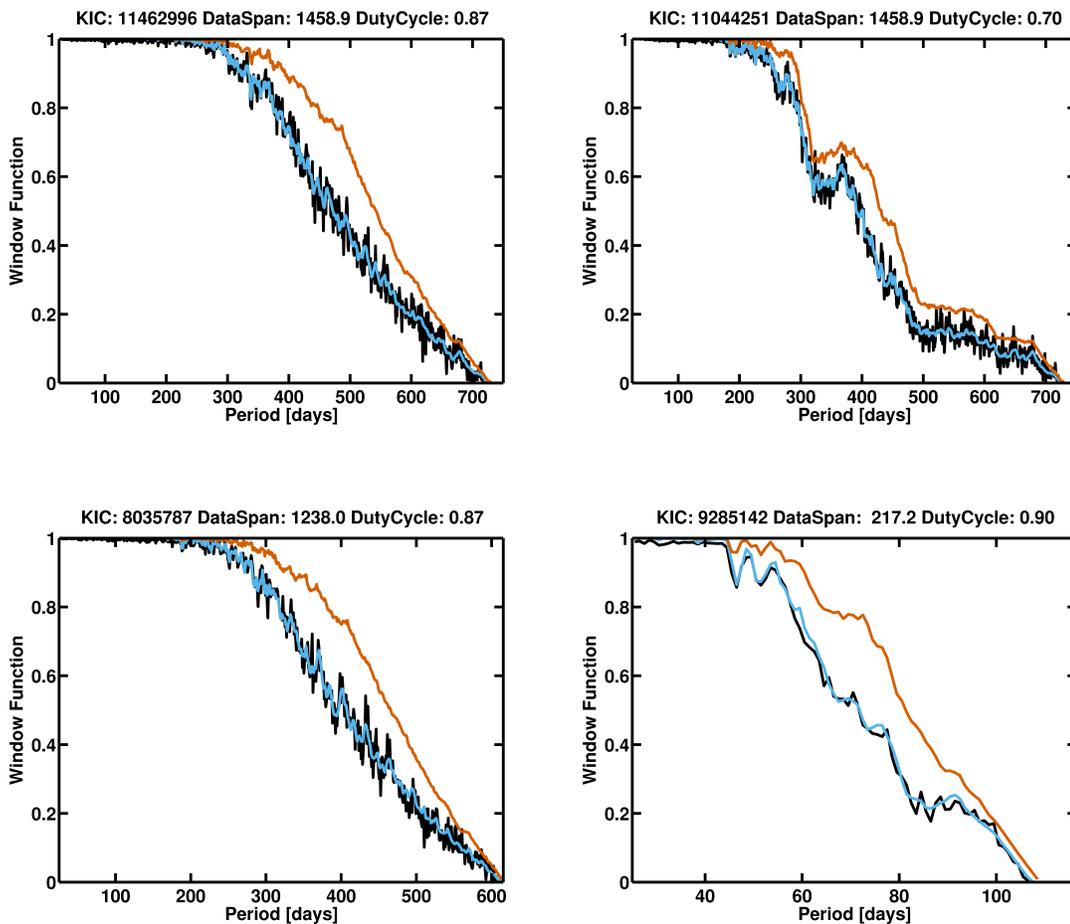


Figure 4. Example window functions of representative targets with varying amounts of observational data. The target's *Kepler* ID, data span, and duty cycle are specified along the top. The previous window function delivery (red line) for DR24 (SOC 9.2) is overly optimistic compared to the empirically derived window function (black line) computed with FLTI using a Monte-Carlo technique. The current window function delivery (DR25; SOC 9.3 – blue line) is in agreement with the empirical window function after including the three-transit weight check.

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